

# A simultaneous *XMM–Newton* and *BeppoSAX* observation of the archetypal broad line Seyfert 1 galaxy NGC 5548

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## ABSTRACT

We report the spectral analysis of a long *XMM–Newton* observation of the well-studied, moderate luminosity broad line Seyfert 1 galaxy NGC 5548. The source was at an historically average brightness and we find the hard (3–10 keV) spectrum can be well fitted by a power law of photon index  $\Gamma \sim 1.75$ , together with reflection. The only feature in the hard X-ray spectrum is a narrow emission line near 6.4 keV, with an equivalent width of  $\sim 60$  eV. The energy and strength of this line are consistent with fluorescence from ‘neutral’ iron distant from the central continuum source. We find no evidence for a broad Fe K line, with an upper limit well below previous reports, suggesting the inner accretion disc is now absent or highly ionized. The addition of simultaneous *BeppoSAX* data allows the analysis to be extended to 200 keV, yielding important constraints on the total reflection. Extrapolation of the hard X-ray power law down to 0.3 keV shows a clear ‘soft excess’ below  $\sim 0.7$  keV. After due allowance for the effects of a complex warm absorber, measured with the *XMM–Newton* RGS, we find that the soft excess is better described as a smooth upward curvature in the continuum flux below  $\sim 2$  keV. The soft excess can be modelled either by Comptonized thermal emission or by enhanced reflection from the surface of a highly ionized disc.

**Key words:** galaxies: active – galaxies: individual: NGC 5548 – galaxies: Seyfert – X-rays: galaxies.

## 1 INTRODUCTION

It is widely believed that the primary luminosity of an active galactic nucleus (AGN) originates in an accretion disc around a supermassive black hole. However, in the X-ray band, a hard power-law component generally dominates above  $\sim 2$  keV in the well studied broad line Seyfert 1 galaxies and is considered to arise in a hot corona above the surface of the accretion disc, where optical/ultraviolet (UV) photons from the disc are Comptonized to X-ray energies. These X-rays in turn illuminate the disc, being either ‘reflected’ towards the observer or thermalized back into optical/UV emission (Mushotzky, Done & Pounds 1993). In the X-ray spectral band, evidence for this disc reflection component, in the form of a fluorescent Fe K $\alpha$  line near 6.4 keV, an Fe K edge at  $>7$  keV and a Compton ‘hump’ at  $>10$  keV, is found in many bright Seyfert 1 galaxies (e.g. Pounds et al. 1990; Nandra & Pounds 1994). The improved resolution of *ASCA* has more recently been used to show that the Fe K $\alpha$  line is often broad, with excess flux, particularly in a ‘red’ wing to the line (Tanaka et al. 1995; Nandra et al. 1997), believed to be due

to emission from the inner disc in a region of both high velocities and high gravity (Fabian et al. 2000).

The substantial gains in sensitivity, bandwidth and spectral resolution of *XMM–Newton* and *Chandra* are now further qualifying this standard picture, with indications that the energy, breadth and strength of the broad Fe K $\alpha$  line (the inner disc component) is critically dependent on how the underlying X-ray continuum is modelled. Also, the frequent detection of a narrow emission line at  $\sim 6.4$  keV, probably having a distinct origin, away from the inner disc (Reeves et al. 2001; Yaqoob et al. 2001; Kaspi et al. 2001; Matt et al. 2001; Gondoin et al. 2001; Pounds et al. 2001), now has to be accounted for in determining the profile of any broad emission line component (Pounds & Reeves 2002; Weaver & Reynolds 1998).

In this paper we report on a simultaneous *XMM–Newton* and *BeppoSAX* observation of NGC 5548, one of the brightest and best-studied, medium luminosity broad line Seyfert 1 galaxies. At  $z = 0.017$  (with  $H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ) NGC 5548 has an X-ray luminosity (2–10 keV) which historically lies in the range  $1\text{--}4 \times 10^{43} \text{ erg s}^{-1}$ . The Galactic absorption column towards NGC 5548 is  $1.65 \times 10^{20} \text{ cm}^{-2}$  (Murphy et al. 1996), rendering it easily visible over the whole ( $\sim 0.2\text{--}12$  keV) spectral band of the EPIC and RGS detectors on *XMM–Newton*.

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The hard X-ray spectrum of NGC 5548 obtained with the *Ginga* satellite (Nandra & Pounds 1994) showed the 2–18 keV spectrum to be well described by a power law of index  $\Gamma \sim 1.65$ – $1.85$ , including a reflection factor  $R = \Omega/2\pi$ , where  $\Omega$  is the solid angle subtended by the reflecting matter, in the range 0.5–1, with an Fe K emission line (equivalent width  $\sim 110$  eV at  $\sim 6.4$  keV) and the imprint of low energy absorption by ionized gas of equivalent hydrogen column density  $5 \times 10^{22}$  atom  $\text{cm}^{-2}$ . Fitting a narrow line to an early *ASCA* observation refined the line energy to  $E = 6.39 \pm 0.04$  keV, with an equivalent width (EW) of  $\sim 90$  eV, increasing to  $\sim 170$  eV for a broad line fit (Nandra et al. 1997).

Evidence for a ‘soft excess’, whereby the X-ray spectrum below  $\sim 1$  keV lies above an extrapolation of the hard X-ray power law, was first implied in spectra of several AGN obtained by *HEAO-1* A2 (Pravdo et al. 1981), *EXOSAT* (Arnaud et al. 1985; Turner & Pounds 1989) and *Einstein* (Bechtold et al. 1987). However, quantifying the soft excess has remained uncertain, as its determination depends on knowledge of both the power-law component and the effects of overlying absorption. The physical origin of any soft excess is also generally unclear, as thermal emission from the accretion disc around a  $10^6$ – $10^8$   $M_\odot$  black hole (e.g. Arnaud et al. 1985; Czerny & Elvis 1987), should be much cooler than the  $\sim 100$  eV soft X-ray temperatures typically inferred, without secondary heating or reprocessing of the disc photons. Understanding the origin of the disc emission, and its interaction with the ‘corona’ nevertheless remains a key to understanding the whole X-ray emission mechanism in Seyfert galaxies.

In the case of NGC 5548, an early *EXOSAT* observation (Branduardi-Raymont 1986) found evidence for a soft excess emerging when the 2–6 keV flux was above some average level. This conclusion gained support from an extended *ROSAT* observation (Done et al. 1995) where the 0.2–0.7 keV flux from NGC 5548 was found to vary more strongly than that at 1–2 keV. However, in a long *BeppoSAX* observation in 1997, Nicastro et al. (2000) found the spectral variability in NGC 5548 was primarily due to a change in the power-law slope, with no requirement for a soft excess. Other results of this *BeppoSAX* study were: a high-energy cut-off in the range 90–165 keV; low-energy absorption well-modelled by O VII and O VIII edges; an unresolved Fe K emission line and a continuum ‘hump’ around 20–30 keV consistent with reflection  $R \sim 0.5$  from ‘cold’ matter.

## 2 OBSERVATION AND DATA REDUCTION

NGC 5548 was observed on 2001 July 9/10 and 12 during orbits 290 and 291 in the *XMM-Newton* Guest Observer programme. The observations with the EPIC MOS2 (Turner et al. 2001) and PN (Struder et al. 2001) detectors were of the order of  $\sim 90$  ks in orbit 290 and  $\sim 35$  ks in orbit 291, with medium and thin filters, respectively. MOS1 data were in timing mode and therefore not useful in the present analysis. The simultaneous RGS observations were of  $\sim 90$  ks and  $\sim 40$  ks, during orbits 290 and 291, while the *BeppoSAX* observation covered the period July 8–10, overlapping the longer *XMM-Newton* (orbit 290) exposure.

*BeppoSAX* data were reduced in the standard way (Guainazzi et al. 1999). MECS counts were extracted from a region of 4 arcmin in radius, and the PDS spectrum was extracted using the fixed pulse rise-time threshold.

The EPIC data were first screened with the XMM SAS v5.3 software. X-ray events corresponding to patterns 0–12 for the MOS2 camera (similar to grades 0–6 in *ASCA*) were used; for the PN patterns 0–4 (single and double pixel events) were selected. A low-

energy cut of 200 eV was applied to all data and known hot or bad pixels were removed. The non X-ray background remained low until the last  $\sim 5$  ks of the orbit 290 observation, which were excluded from the subsequent spectral analysis.

We extracted source and background spectra for the PN and MOS2 detectors with a circular source region of 45 arcsec in radius defined around the centroid position of NGC 5548, with the background taken from an offset position close to the source. The 2–10 keV X-ray flux from NGC 5548 varied between  $5$ – $7 \times 10^{-11}$  erg  $\text{cm}^{-2}$   $\text{s}^{-1}$  during the *XMM-Newton* observations, lying at the upper end (though near the mean) of the historical range; however, photon pile-up was negligible in the Small Window mode chosen for all EPIC observations. The individual spectra were binned to a minimum of 20 counts per bin, to facilitate use of the  $\chi^2$  minimalization technique in spectral fitting.

The primary analysis in this paper is based on observations from the first *XMM-Newton* orbit, which has the longest exposure and also benefits from simultaneous *BeppoSAX* data. In the present analysis we use both MOS2 and PN data to evaluate the hard X-ray spectrum, but only the MOS2 data to study the ‘soft excess’, since the MOS2 broad-band calibration is considered to be more certain than that of the PN at present. Response functions for spectral fitting to the *XMM-Newton* data were generated from SAS v5.3. We use the MECS and PDS data from *BeppoSAX* to constrain the continuum reflection and high energy cut-off.

All spectral fits include absorption due to the line-of-sight Galactic column of  $N_H = 1.65 \times 10^{20}$   $\text{cm}^{-2}$ . All fit parameters are given in the AGN rest-frame. Errors are quoted at the 90 per cent confidence level (e.g.  $\Delta\chi^2 = 2.7$  for one interesting parameter).

## 3 THE HARD POWER-LAW SPECTRUM

It has long been clear that the key to modelling the X-ray spectra of AGN is to first identify, if possible, the ‘underlying continuum’ on which additional emission and absorption features are superimposed. We therefore began our analysis of NGC 5548 by fitting the *XMM-Newton* data above 3 keV, where any effects of intervening absorption should be negligible, a view supported by the RGS analysis of the warm absorber (Steenbrugge et al. 2003). A power law gave a good fit to both MOS2 and PN data over the 3–10 keV band, apart from a clear excess flux near 6.4 keV. The respective parameters are listed in fit 1 of Tables 1 and 2. The power-law index was flatter by  $\sim 0.02$  in the PN fit, a difference that is now only marginally significant and suggests an encouraging convergence in the latest calibrations of the EPIC cameras. Fig. 1 reproduces the data: model ratio of the MOS2 and PN power-law fits. No other features are seen in addition to the narrow emission line at  $\sim 6.4$  keV.

### 3.1 The Fe K band

The addition of a narrow Gaussian emission line to the simple power-law fits, with energy, width and flux free, produced a very significant improvement to both MOS2 and PN spectral fits. These improved fits are detailed in fit 2 of Tables 1 and 2, respectively. Inclusion of the emission line in the 3–10 keV fits increased the power-law index by  $\sim 0.03$  for each data set. The line is narrow and lies at a rest energy consistent with fluorescence from neutral iron. More specifically, the upper limit of 6.41 keV in both MOS2 and PN fits corresponds to ion stages between Fe I and Fe X, while it is interesting to note that the high statistical significance of the EPIC data gives a linewidth with similar formal precision to that reported from an earlier observation

**Table 1.** Spectral fits to joint MOS, MECS and PDS data over 3–10 keV and 3–200 keV energy bands.

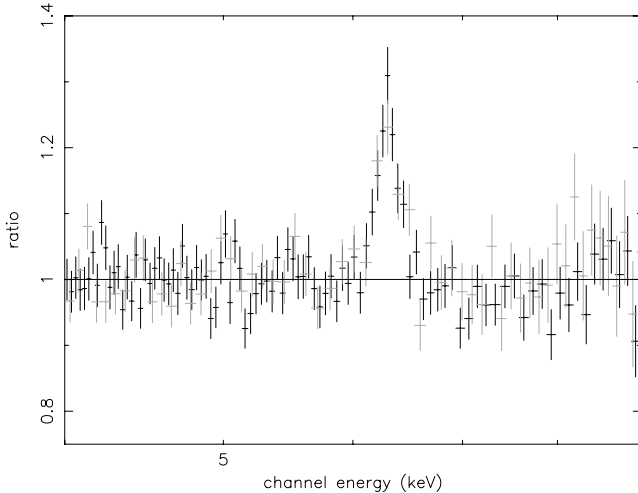
Fit	$\Gamma$	Fe K line			Reflection component		$\chi^2/\text{dof}$
		$E^a$	$\sigma^b$	EW <sup>c</sup>	$E_c^d$	$R^e$	
1. P-L only	$1.67 \pm 0.02$						426/399
2. P-L + Gauss	$1.70 \pm 0.02$	$6.39 \pm 0.02$	$40 \pm 40$	$63 \pm 18$			354/396
3. P-L + Gauss + Pextrav	$1.75 \pm 0.02$	$6.39 \pm 0.02$	$45 \pm 45$	$63 \pm 18$	$240 \pm 90$	$0.75 \pm 0.19$	408/467

Notes. <sup>a</sup>Rest energy of line (keV). <sup>b</sup>Intrinsic (1 sigma) width of line (eV). <sup>c</sup>Equivalent width of line (eV). <sup>d</sup>Energy cut-off (keV). <sup>e</sup>Reflection parameter of cold matter;  $R = \Omega/2\pi$ .

**Table 2.** Spectral fits to joint PN, MECS and PDS data over 3–10 keV and 3–200 keV energy bands.

Fit	$\Gamma$	Fe K line			Reflection component		$\chi^2/\text{dof}$
		$E^a$	$\sigma^b$	EW <sup>c</sup>	$E_c^d$	$R^e$	
1. P-L only	$1.65 \pm 0.01$						1511/1353
2. P-L + Gauss	$1.68 \pm 0.01$	$6.40 \pm 0.01$	$64 \pm 24$	$67 \pm 14$			1347/1350
3. P-L + Gauss + Pextrav	$1.73 \pm 0.02$	$6.39 \pm 0.02$	$60 \pm 24$	$60 \pm 18$	$275 \pm 100$	$0.58 \pm 0.15$	1440/1465

Notes. <sup>a</sup>Rest energy of line (keV). <sup>b</sup>Intrinsic (1 sigma) width of line (eV). <sup>c</sup>Equivalent width of line (eV). <sup>d</sup>Energy cut-off (keV). <sup>e</sup>Reflection parameter of cold matter;  $R = \Omega/2\pi$ .



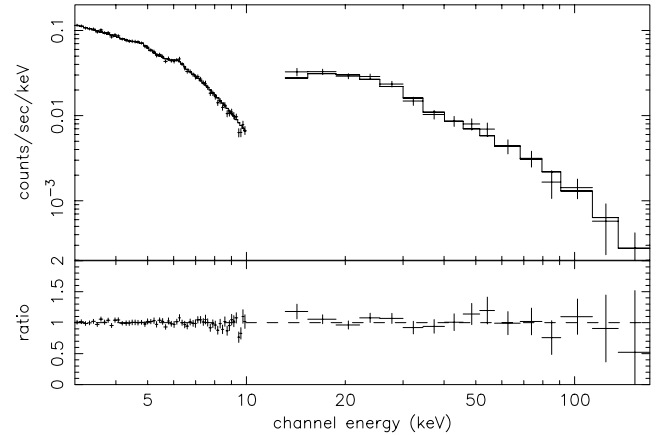
**Figure 1.** MOS2 (black) and PN (grey) spectral fits over the 3–10 keV band shown as the ratio of the data to the best power-law fit in each case. The only spectral feature evident is a narrow emission line at a rest energy of  $\sim 6.4$  keV. This figure can be seen in colour in the online version of the journal on *Synergy*.

with the *Chandra* HETG (Yaqoob et al. 2001). We return to this point in the Discussion (Section 5).

The formal  $3\sigma$  upper limit for an additional broad line, modelled as a ‘diskline’ (Fabian et al. 1989) with energy and profile as reported from an *ASCA* observation (Nandra et al. 1997) is 43 eV. In contrast, the *ASCA* analysis, which recorded a power-law index and flux almost identical to the present observation, found a broad line equivalent width of  $200 \pm 90$  eV.

### 3.2 Cold reflection

An important limitation in *XMM–Newton* data is a rapid loss of sensitivity above  $\sim 10$  keV, preventing the unambiguous determination of continuum reflection, which is in general required to determine the true slope of the primary (power-law) continuum. Fortunately we have simultaneous *BeppoSAX* data on NGC 5548 for the longer, orbit



**Figure 2.** Combined MECS and PDS spectral fit over the 3–200 keV range compared to the cutoff power law plus cold reflection model of Table 1, fit 3. The upper panel shows the count spectra (crosses) and the model folded through the detector responses (histograms). The lower panel shows the fit residuals.

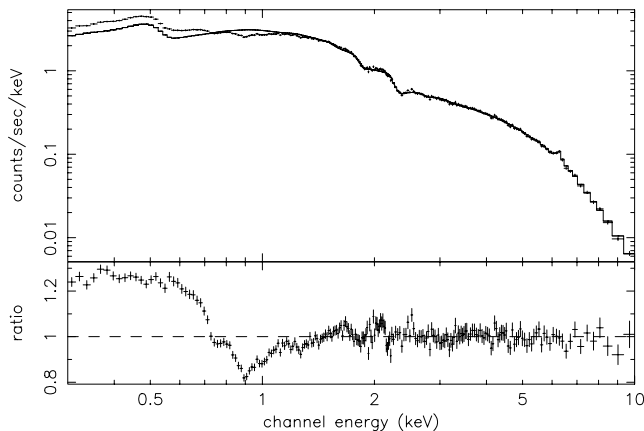
290 observation, which extends the spectral coverage to  $\sim 200$  keV and constrains the continuum ‘hump’ caused by down-scattering of high-energy photons and is a key measure of reflection. We first compared the MOS2, PN and MECS data, finding good agreement in describing the NGC 5548 spectrum over the band 3–10 keV. The PDS and MECS data were then combined, and tested against a model with components of a cut-off power law, a narrow Fe K line and cold reflection modelled by PEXRAV. All fit parameters were tied for the two detectors, with an additional normalization constant of 0.86 for the PDS (Fiore, Guainazzi & Grandi 1999). An excellent fit (Fig. 2) yielded a cold reflection factor  $R \sim 0.7$  which, in turn, led to a steepening of the power-law index by  $\sim 0.05$ . Addition of the MOS2 data, with again only the normalization untied, confirmed the excellence of this fit, details of which are listed in Table 1, fit 3. The parameters of the reflection fit were found to be only a weak function of the assumed iron abundance. In the best fit, quoted, the abundance was left free and determined to be

$1.4 \pm 0.6$  solar. The cosine of the disc inclination was fixed at 0.9. The PN, MECS and PDS data were next fitted with the above model, yielding the parameters detailed in Table 2, fit 3.

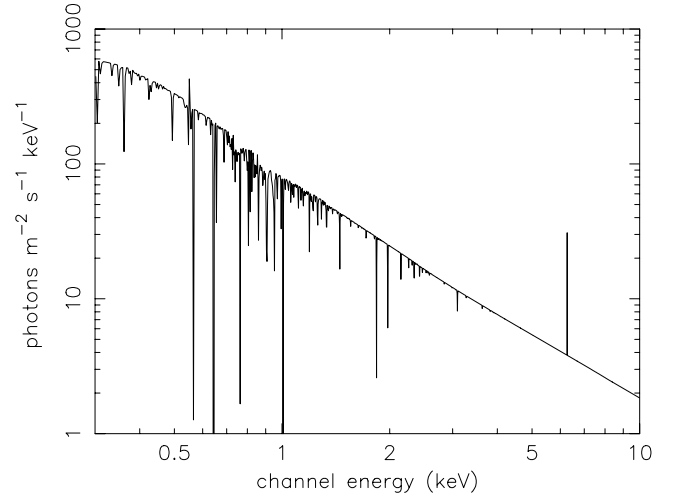
In summary, an overall and self-consistent description of the 3–200 keV spectrum consists of a cut-off power law and cold reflection at a level that could also explain the narrow Fe K emission line (not included in PEXRAV). The observed 3–100 keV flux was  $1.2 \times 10^{-10}$  erg cm $^{-2}$  s $^{-1}$ .

#### 4 THE SOFT EXCESS

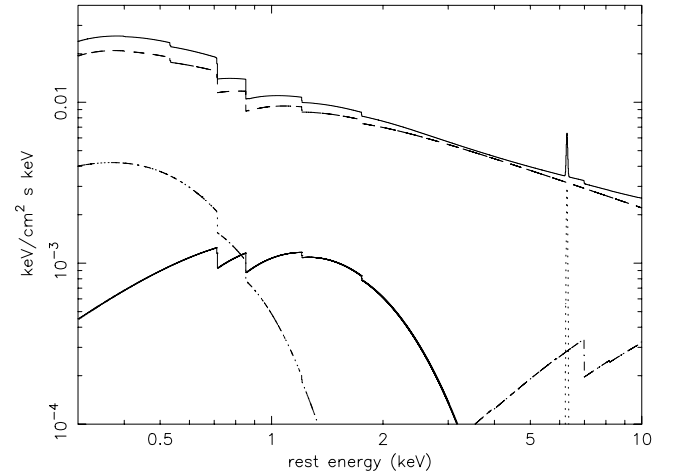
Extrapolating the best-fitting description of the hard X-ray spectrum (fit 3 in Table 1) to the lower well-calibrated limit of the EPIC MOS camera ( $\sim 0.3$  keV) shows an apparently unambiguous excess of soft X-ray flux below  $\sim 0.7$  keV (Fig. 3). There is also clear evidence for absorption which ‘cuts’ a significant fraction from any smooth profile continuum between  $\sim 0.7$ –2 keV, a factor that appears to be common in the EPIC spectra of low- and medium-luminosity AGN and must be accounted for before attempting to model any soft excess emission (Pounds & Reeves 2002). In the present case the simultaneous *XMM-Newton* RGS data (Steenbrugge et al. 2003) confirm the presence of complex absorption and associate it with an ionized outflow. We have worked with colleagues at Space Research Organisation of the Netherlands (SRON) to assess joint fits to the EPIC and RGS data for NGC 5548 (see Fig. 4). Our conclusions are that the data are in good overall agreement, though the RGS data are relatively insensitive to the precise continuum shape, which in the joint fit is effectively determined by the EPIC data. For the present EPIC analysis a simpler description of the warm absorber was used, with advantages in terms of modelling speed and in being a better match to the resolution of the EPIC detectors. This working description of the warm absorber is in terms of a series of absorption edges at (rest frame)  $\sim 0.725$  keV ( $\tau \sim 0.29$ ),  $\sim 0.87$  keV ( $\tau \sim 0.28$ ),  $\sim 1.23$  keV ( $\tau \sim 0.08$ ) and  $\sim 1.78$  keV ( $\tau \sim 0.06$ ). The energy and optical depth of these edges were left free in fitting to the EPIC data, and correspond approximately to the absorption edges of O VII, O VIII, Ne IX/X and Mg XI/XII. Visual comparison of Figs 4 and 5 indicates why the absorption edge set is a reasonable substitute for the complex RGS absorption spectrum at the resolution level of the EPIC detectors. We note that the two deepest edges are consistent with the main features found in *ROSAT* and *ASCA* modelling, and account for the strongest absorption, from O VII, O VIII and iron M- and L-shell ions, while Ne K-shell absorption is also clearly seen.



**Figure 3.** Hard X-ray spectral fit above 3 keV, fit 3 of Table 1, extrapolated to 0.3 keV showing the soft excess modified by line-of sight absorption (see text for details).



**Figure 4.** Joint MOS2 and RGS data fit for NGC 5548 over the 0.3–10 keV band. The underlying continuum is modelled with a power law and two blackbodies (as in Table 3, fit 1), with complex absorption determined by the RGS.



**Figure 5.** Model fit to the MOS2 data over the 0.3–10 keV range with a power law plus two blackbody emitters as detailed in Table 3, fit 1. The warm absorber is modelled here with a set of four absorption edges as described in the text.

Fixing the cold reflection factor and Fe K line parameters as determined from the 3–200 keV fit, we then proceeded to model the ‘soft excess’. A second power law, or a single blackbody, gave a very poor fit, the breadth of the soft excess requiring two blackbody components to obtain an acceptable value of chi-squared. Table 3 (fit 1) details this best fit to the overall (0.3–10 keV) spectrum, which is reproduced in Fig. 5. Parametrized in this way, the soft excess is found to be energetically a minor part of the X-ray flux from NGC 5548, the multiple blackbody component accounting for only  $\sim 20$  per cent of the observed 0.3–1 keV flux of  $2.1 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ ,  $\sim 12$  per cent of the 1–2 keV flux of  $1.5 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ , and a negligible fraction of the harder (2–10 keV) flux of  $5.0 \times 10^{-11}$  erg cm $^{-2}$  s $^{-1}$ .

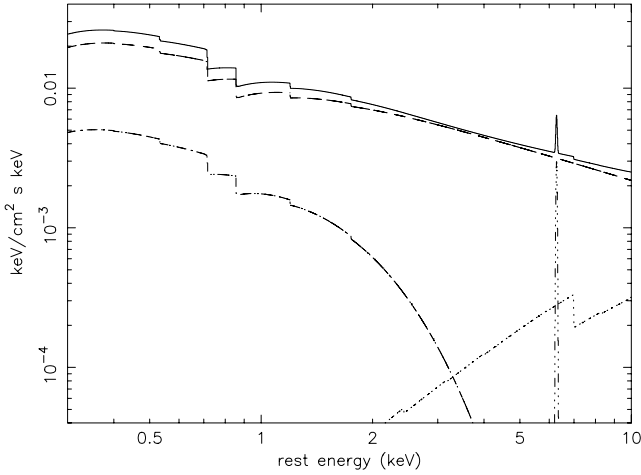
#### 4.1 Comptonized disc emission

While the addition of multiple blackbody components is a useful way to parametrize the soft excess it reveals little about the physical

**Table 3.** Broad-band (0.3–10 keV) continuum spectral fits to MOS2 data.

Fit	$\Gamma^a$	$kT_1^b$	$kT_2^b$	$\tau^c$	$\xi^d$	$\chi^2/\text{dof}$
1. PL+2 $\times$ bbody	$1.76 \pm 0.03$	$130 \pm 8$	$365 \pm 19$			551/566
2. PL+compTT	$1.75 \pm 0.03$		$400 \pm 60$	$16 \pm 4$		581/569
3. PL+DB <sup>f</sup>	$1.75 \pm 0.02$				$4.1 \pm 0.2$	586/569

Notes. <sup>a</sup>Photon Index of power-law. <sup>b</sup>Temperature (eV) of blackbody emission or Comptonizing electrons. <sup>c</sup>Optical depth of Comptonizing electron distribution. <sup>d</sup>Ionization parameter. <sup>e</sup>DB = Fit to ionized reflection model (Ballantyne et al. 2001) – see text for details.



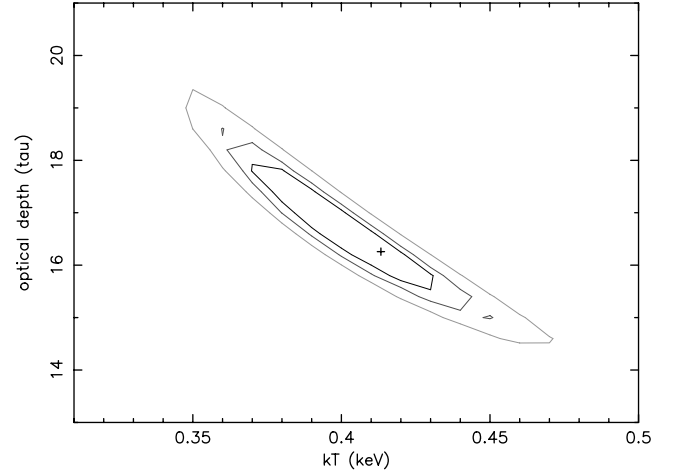
**Figure 6.** Model fit to the MOS2 data over the 0.3–10 keV range with a power law plus Comptonized thermal emission as detailed in Table 3, fit 2. The warm absorber is again modelled with a set of 4 absorption edges.

origin of this emission component, because – as noted earlier – the intrinsic thermal radiation of an accretion disc around a supermassive black hole is expected to be much cooler than implied by the observed soft X-ray spectrum.

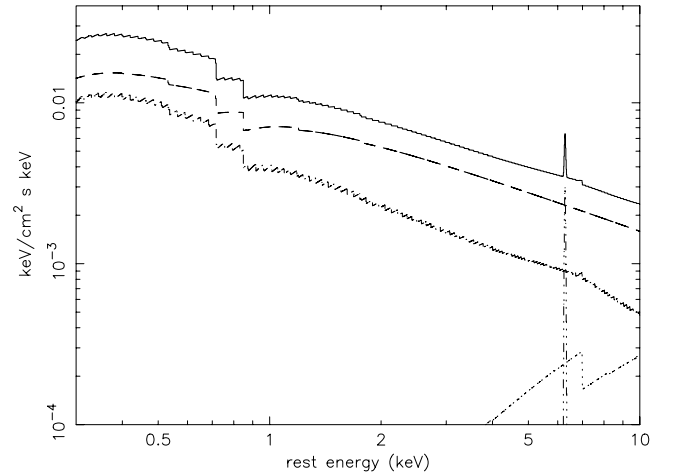
To attempt a more physical description of the spectral curvature underlying the complex absorption, we therefore replaced the blackbody components with Comptonized thermal disc emission, a model successfully used to describe *XMM–Newton* detections of strong soft excesses in PKS 0558–504 (O’Brien et al. 2001) and 1H 0419–577 (Page et al. 2002), and applied to *BeppoSAX* data of NGC 5548 (Petrucci et al. 2000). Specifically, we used the *COMP*TT model in *XSPEC* which gave a good fit to the 0.3–10 keV MOS spectrum of NGC 5548, with parameters as detailed in Table 3, fit 2 and illustrated in Fig. 6. We assumed an input photon temperature of 25 eV, and found an optically thick scattering region ( $\tau \sim 16$ ) of  $kT \sim 0.4$  keV. These parameters were found to be relatively insensitive to the input photon temperature over a factor of 2–3. As they are also known to be strongly covariant, we show in Fig. 7 the range of acceptable values for the temperature and optical depth. We interpret the relatively low temperature of the scattering medium (compared with previous fits to the soft excess of PKS 0558–504 and 1H 0419–577) in terms of matching the less broad (less hot) soft excess in NGC 5548. An important prediction of this fit is that soft X-ray variability would be reduced by the large optical depth (in contrast to the conclusions of Done et al. 1995).

#### 4.2 Ionized reflection

We next attempted a fit with an ionized reflector model (e.g. Ballantyne, Ross & Fabian 2001). In this model, when the disc



**Figure 7.** 1, 2 and 3 $\sigma$  contours of the covariant parameters of temperature and optical depth in the Comptonized model fit to the soft excess. This figure can be seen in colour in the online version of the journal on *Synergy*.



**Figure 8.** Model fit to the MOS2 data over the 0.3–10 keV range with a power law plus ionized reflection (lower curve) as detailed in Table 3, line 3. The warm absorber is again modelled with a set of four absorption edges.

surface is highly ionized it produces a smooth excess of emission in the soft band, with relatively weak spectral features. The result was again a good fit to the soft X-ray excess, with an ionized reflection factor of  $R \sim 0.6$ . Other details of this fit are given in Table 3, fit 3, and the model is reproduced in Fig. 8.

At the high ionization parameter implied by this fit, individual spectral features are smeared out, due to the occurrence of multiple ion stages and to the effects of electron scattering in the upper layers

of the reflector. This smearing is critical in fitting to our NGC 5548 observation over the  $\sim 7\text{--}9$  keV band where deep absorption edges of ionized Fe would otherwise be in conflict with the data.

## 5 DISCUSSION

Our analysis emphasizes the importance of simultaneous broad-band data to de-convolve the several inter-related components in the observed X-ray spectrum of AGN. Without such data, progress in understanding the primary emission and re-processing mechanisms will remain elusive. The key importance of measuring continuum reflection is underlined in our study of NGC 5548 where it has a significant effect on the deduced hard X-ray power-law slope, which in turn is critical in quantifying any soft excess component. The impact of overlying absorption on the emerging emission in the soft X-ray band must also be properly accounted for to determine the true soft emission spectrum.

With the above constraints in mind, our analysis of the long observation of NGC 5548 on 2001 July 9–10 proceeded by first determining the amount of cold reflection using the simultaneous *BeppoSAX* data. The value obtained of  $R \sim 0.58\text{--}0.75$  for the *BeppoSAX* plus PN and *BeppoSAX* plus MOS2 fits is compatible with the strength of the narrow Fe K line, from which one conclusion could be that the only significant reflection is from cold matter some distance from the central continuum source. That would imply that the inner accretion disc is either absent or highly ionized, a conclusion consistent with the absence of a broad Fe K line.

In the latter respect it is interesting that one description of the (relatively weak) soft excess is by enhanced reflection from highly ionized matter, presumably the inner accretion disc. However this interpretation of the soft excess is not unique, and we have no useful constraints in terms of either ionized absorption edges in the  $\sim 8\text{--}9$  keV region (where the *XMM-Newton* sensitivity is rather low) or a definite contribution to the continuum ‘hump’ at higher energies. The latter point is considered further in Section 5.3.

### 5.1 Narrow Fe K line

The best prospect for clarifying the origin of the narrow Fe K line, with current observational facilities, appears to be via a careful search for variability with the statistical precision of *XMM-Newton* EPIC data, or by constraining the linewidth with the HETG on *Chandra*. A comparison of our present data with the *Chandra* observation of Yaqoob et al. (2001) provides a basis for such a future study in the case of NGC 5548.

Taking our PN data, which has the best statistics on the iron line, we find a line energy of  $6.40 \pm 0.01$  keV, consistent with the *Chandra* line energy of  $6.402 \pm 0.026$  keV. The line fluxes are also consistent, albeit not well constrained. Given current uncertainties in cross-calibration between the two observations, we note instead that the continuum flux was a factor of 2 lower during the *Chandra* observation, while the EW was double the *XMM-Newton* value (Yaqoob et al. 2001 find  $133 \pm 58$  eV, compared with our *XMM-Newton* figure of  $60 \pm 18$  eV), suggesting that the bulk of the narrow line may have remained constant over the 17-month interval between the *Chandra* and *XMM-Newton* observations.

However, we also note that the PN data marginally resolve the ‘narrow’ line, with  $\sigma = 60 \pm 24$  eV ( $6500 \pm 2600$  km s $^{-1}$  full width at half maximum, hereafter FWHM). The comparable *Chandra* HETG parameters were  $4500^{+3560}_{-2590}$  km s $^{-1}$  FWHM. Effectively, the much better statistics of the *XMM-Newton* detection are compensating for the fourfold worse energy resolution, lending some

support to the conclusion that the ‘narrow’ Fe K line in NGC 5548 is marginally resolved. If confirmed, the implication of this result is that a significant component of the ‘narrow’ Fe K line arises at distances from the hard X-ray source comparable to the broad line clouds.

### 5.2 Soft excess

Having obtained a reliable measure of the dominant power-law component, extrapolation to the lower useful energy limit of EPIC confirms the presence of a ‘soft excess’. Given the simultaneous measure of the ‘warm absorber’ by the *XMM-Newton* RGS, we feel confident that – in this case – the soft excess is reliably determined. We find it to be characterized by a smooth upward curvature, probably typical of unobscured, radio-quiet AGN (Pounds & Reeves 2002), which can be parametrized by blackbody emitters of  $\sim 100$  and  $\sim 300$  keV.

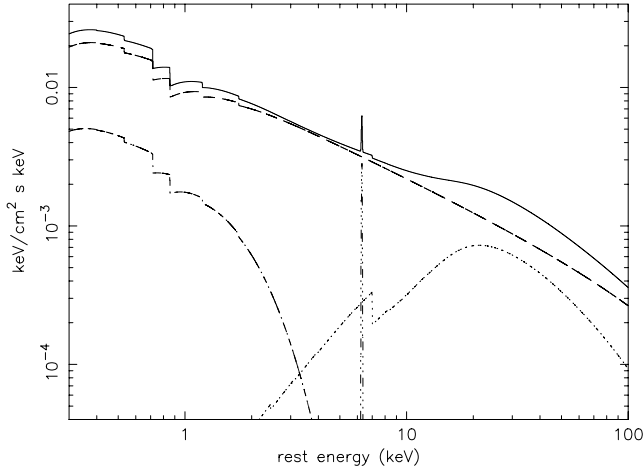
An explanation of the soft excess remains unclear, however. Two physical models, which we find to be equally good fits to the EPIC data, have different inferences for the X-ray emission process. Ionized reflection is an attractive explanation, given the energetically dominant and relatively hard power law spectral component in NGC 5548. This implies that in NGC 5548 – which we point out may be an archetypal broad line Seyfert 1 galaxy – the X-ray emission is primarily due to accretion energy being deposited in a low density, high temperature medium (eg a ‘corona’), and there is little or no intrinsic X-ray emission from the accretion disc.

On the other hand, if the soft excess arises by thermal disc emission, up-scattered into the soft X-ray band by passage through an optically thick, hot ‘skin’ (e.g. Hubeny et al. 2001), it offers the important possibility of studying the thermal structure and internal heating of the inner disc. Some support is provided for the Comptonization model by the finding of a time-lag between extreme UV (EUV) and X-ray flux variations from NGC 5548 in simultaneous observations with *ASCA*, *RXTE* and *EUVE* (Chiang et al. 2000).

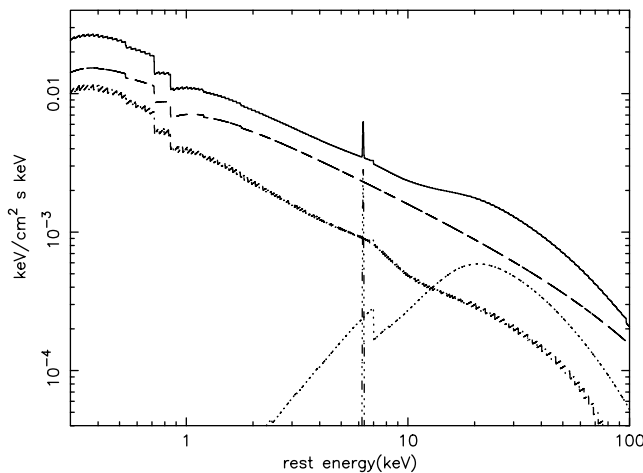
In principle, it should be possible to distinguish between the above two models in the higher energy spectral band where the ionized reflector will continue to make a contribution. We have attempted this comparison by extending the Comptonized thermal and ionized reflector fits to 100 keV by inclusion of the PDS data. The power-law index was fixed in each case at the MOS2 fit value of  $\Gamma = 1.75$ , and the warm absorber and Fe K emission line parameters were also retained from the 0.3–10 keV fits. Good fits were obtained with both models, which yielded interesting (but not decisive) differences in 10–100 keV band. The Comptonized thermal description of the soft excess ( $\chi^2/\text{dof} = 576/590$ ) has essentially no effect in the higher energy band (Fig. 9), with all the continuum ‘hump’ being due to the cold reflector. In the case of the ionized reflector fit (only marginally worse with  $\chi^2/\text{dof} = 593/590$ ), there is a contribution across the whole observed band (Fig. 10). In consequence, the Fe abundance in PEXRAV drops from  $\sim 1.3$  to  $\sim 0.3$  to compensate for the ‘filling in’ of the observed flux at  $\sim 7\text{--}15$  keV. In addition the high energy cut-off energy of  $\sim 150$  keV in the Comptonization fit extends beyond  $\sim 300$  keV in the ionized disc fit. We conclude that, in the present case, our high energy data are unable to distinguish between the two models for the (relatively weak) soft excess.

### 5.3 The X-ray signature of NGC 5548, a low accretion rate AGN

The average observed flux in the 2–10 keV band during orbit 290 was  $5 \times 10^{-11}$  erg s $^{-1}$  cm $^{-2}$ , corresponding to a luminosity of  $\sim 2.5 \times$



**Figure 9.** Power-law, cold reflection and Comptonized thermal emission model fit to the MOS2 and PDS data. See text for further details.



**Figure 10.** Power-law, cold reflection and ionized reflection fit to the MOS2 and PDS data. The latter is seen to make only a small contribution to the high energy ‘hump’. See text for further details.

$10^{43}$  erg  $s^{-1}$ . The bolometric luminosity is then likely to be of order  $\sim 5 \times 10^{44}$  erg  $s^{-1}$ . Assuming a black hole mass for NGC 5548 of  $M \sim 6 \times 10^7 M_{\odot}$  (Ferrarese et al. 2001) this luminosity requires an accretion rate of  $\sim 4$  per cent of the Eddington rate. We suggest this may be a typical value for many broad line Seyfert 1 galaxies.

The corresponding X-ray signature of broad line Seyfert 1s might then be much as we find for NGC 5548, namely a dominant power-law component with (hard) photon index  $\sim 1.75$ , a weak but relatively ‘hot’ soft excess, and a narrow, ‘cold’ Fe K emission line. The most clear-cut reflection component, dominating the  $\sim 10$ – $100$  keV spectrum of NGC 5548, appears also to be from cold (perhaps the same) matter subtending  $\sim 1$ – $1.5\pi$  sr. The emergence of a similar value in Compton-thick Seyfert 2 galaxies (Matt 2002; Schurch 2002) suggests that substantial reflection from cold matter, distant from the hard X-ray source, is a common property of all Seyfert galaxies. If so, this underlines the importance of taking due account of such reflection in any analysis or modelling of AGN spectra.

The absence of a broad iron K line is consistent with the inner accretion disc being highly ionized by the relatively hard incident spectrum, or simply absent at the low implied accretion rate. We

note that the former explanation would be qualitatively consistent with one of the two physical models fitted to the soft excess, with magnetic flares providing the localized source(s) of hard X-ray flux.

In the context of the alternative, thermal Comptonization description of the soft excess, it is interesting to recall that *EXOSAT* data showed that the ‘soft excess’ in NGC 5548 was only apparent when the source flux exceeded  $3.5 \times 10^{-11}$  erg  $s^{-1}$   $cm^{-2}$  (Branduardi-Raymont 1986), as it is here. We note that a flux-related ‘switching on’ of the disc X-ray emission in AGN could be a further analogy with the X-ray properties of Galactic black hole candidate sources, which are found to change from a faint/hard state when the accretion rate rises above  $\sim$  a few per cent of the Eddington rate (e.g. Chen et al. 1995). A common physical explanation might then be sought in terms of the well known radiation pressure instability (Shapiro, Lightman & Eardley 1976), with re-structuring the inner disc leading to Comptonization of the thermal disc photons at a relative accretion rate of a few per cent (Kubota, Makashima & Ebisawa 2001; Done 2002).

A truncated disc provides the most obvious alternative (to the highly ionized inner disc), consistent with the X-ray signature of NGC 5548, namely a hard power-law spectrum, moderate luminosity, weak soft excess, and an absence of a broad iron K line. In this scenario the inner radius of the dense matter should penetrate further into a hot inner flow as the accretion rate increased, perhaps giving a more gradual spectral change.

Future studies of the broad-band spectral variability of sources such as NGC 5548 will be important to search for such effects, as will the testing of models over a wider sample of AGN, with a larger range of spectral properties.

## 6 CONCLUSIONS

(i) A long *XMM-Newton* and *BeppoSAX* observation of the bright broad line Seyfert 1 galaxy finds a broad-band X-ray spectrum dominated by a hard power law of index  $\Gamma \sim 1.75$ , a value derived after due allowance for a continuum reflection component peaking at  $\sim 30$ – $40$  keV.

(ii) The continuum reflection corresponds to cold (neutral) matter subtending a solid angle of  $\sim 1$ – $1.5\pi$  sr at the hard X-ray source. We suggest this may be a very common component in the nuclei of Seyfert galaxies.

(iii) The only spectral feature seen in the 3–10 keV spectrum is a narrow emission line at  $\sim 6.4$  keV with an equivalent width of  $\sim 60$  eV, consistent with iron fluorescence from the same cold matter. The lack of variability over a 17 month interval between *Chandra* and *XMM-Newton* observations is consistent with this matter being in the putative torus; however, the marginally resolved linewidth suggests a significant component lies at a smaller radius.

(iv) A weak soft excess is seen, after allowing for over-lying absorption, as a smooth upward curvature in the X-ray continuum below  $\sim 2$  keV, and can be modelled by two blackbodies, Comptonized thermal emission from the accretion disc, or ionized disc reflection.

(v) The overall X-ray emission spectrum is consistent with a model in which accretion at a relatively low rate (a few per cent of the Eddington rate) is depositing most of the corresponding accretion energy in a hot, low-density medium (or corona), with the dense inner disc being absent (a truncated disc) or highly ionized.

(vi) The  $3\sigma$  upper limit of 43 eV on a broad iron K line confirms the inner disc to be in a very different state from that indicated by the earlier *ASCA* observation, although the hard X-ray power law emission was essentially the same.

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